

Strong Evidence for a Buried AGN in UGC 5101: Implications for LINER-Type Ultra-Luminous Infrared Galaxies

Masatoshi Imanishi¹

National Astronomical Observatory, Mitaka, Tokyo 181-8588, Japan

imanishi@optik.mtk.nao.ac.jp

C. C. Dudley²

Naval Research Laboratory, Remote Sensing Division, Code 7217, Building 2, Room 240B, 4555 Overlook Ave SW, Washington DC 20375-5351, U.S.A.

dudley@vivaldi.nrl.navy.mil

and

Philip R. Maloney

Center for Astrophysics and Space Astronomy, University of Colorado, Boulder, CO 80309-0839, U.S.A.

maloney@casasrv.colorado.edu

ABSTRACT

We report on the results of 3–4 μm spectroscopy of the ultra-luminous infrared galaxy (ULIRG) UGC 5101. It has a cool far-infrared color and a LINER-type optical spectrum, and so, based on a view gaining some currency, would be regarded as dominated by star formation. However, we find that it has strong 3.4 μm carbonaceous dust absorption, low-equivalent-width 3.3 μm polycyclic aromatic hydrocarbon (PAH) emission, and a small 3.3 μm PAH to far-infrared luminosity ratio. This favors an alternative scenario, in which an energetically dominant AGN is present behind obscuring dust. The AGN is plausibly obscured along all lines of sight (a ‘buried AGN’), rather than merely obscured along our particular line of sight. Such buried AGNs have previously been found in thermal infrared studies of the ULIRGs IRAS 08572+3915 and

¹Visiting Astronomer at the Infrared Telescope Facility, which is operated by the University of Hawaii under contract from the National Aeronautics and Space Administration.

²NRL/NRC Research Associate

IRAS F00183–7111, both classified optically as LINERs. We argue that buried AGNs can produce LINER-type optical spectra, and that at least some fraction of LINER-type ULIRGs are predominantly powered by buried AGNs.

Subject headings: galaxies: active — galaxies: nuclei — infrared: galaxies – galaxies: individual (UGC 5101)

1. Introduction

Ultra-luminous infrared galaxies (ULIRGs) radiate most of their extremely large, quasar-like luminosities ($> 10^{12} L_{\odot}$) as infrared dust emission, and dominate the bright end of the galaxy luminosity function in the nearby universe (Soifer et al. 1987). Recent studies have revealed that the bulk of the cosmic sub-mm background emission has been resolved into discrete sources, similar to nearby ULIRGs (Blain et al. 1999a), and for this reason data on nearby ULIRGs have been extensively used to derive information on star-formation rates, dust content, and metallicity in the early universe (Barger, Cowie, & Richards 2000; Blain et al. 1999b). Understanding the nature of nearby ULIRGs is therefore of particular importance both locally and cosmologically.

In contrast to the majority of less infrared luminous ($< 10^{12} L_{\odot}$) galaxies, it has been suggested that energetically important, dust-obscured AGNs are present in ULIRGs (Veilleux, Kim, & Sanders 1999a; Fischer 2000). If the geometry of the dust distribution is toroidal and/or the amount of dust along our line of sight is not too great, signatures of such dust-obscured AGNs can be found in the optical and/or near-infrared wavelength range at $\lambda < 2 \mu\text{m}$ (Veilleux et al. 1999a; Veilleux, Sanders, & Kim 1999b). However, detecting AGNs that are deeply embedded in a spherical dust shell (hereafter buried AGNs) and estimating their energetic importance is very difficult at $\lambda < 2 \mu\text{m}$, even though they are energetically important, since spectral tracers are dominated by less strongly obscured star-formation activity. To detect such buried AGNs and estimate their energetic importance, we must observe at wavelengths where the effects of dust extinction are smaller.

At wavelengths 3–13 μm , dust extinction is lower. Furthermore, emission is dominated by radiation from dust, the dominant emission mechanism of ULIRGs. Thus, discussions of the energy sources of ULIRGs are more straightforward than those based on observations at X-ray or radio wavelengths, which are also potentially capable of detecting buried AGNs, but are dominated by emission mechanisms other than dust emission. Finally, by using spectral features at 3–13 μm , we can distinguish the energy sources of individual galaxies: while star-formation-dominated galaxies show strong polycyclic aromatic hydrocarbon (PAH) emission features, buried AGNs should show dust absorption features below smooth continuum emission (Roche et al. 1991; Genzel et al. 1998; Dudley 1999; Imanishi & Dudley 2000). Using spectral features, particularly the 7.7 μm PAH emission, in the rest-frame 5–11 μm spectra obtained with the *Infrared Space Observatory* (ISO), it has been argued that the majority of ULIRGs are star-formation powered (Genzel et al. 1998; Rigopoulou et al. 1999; Tran et al. 2001). However, the determination of the continuum level

with respect to which emission and absorption features should be measured in *ISO* spectra is highly uncertain, due to insufficient wavelength coverage longward of these features. Distinguishing between star-formation and buried AGN activity therefore remains controversial (Dudley 1999; Spoon et al. 2001).

As discussed in Imanishi & Dudley (2000), observations at 3–4 μm have two important advantages: firstly, dust extinction is as small as that at 7–8 μm (Lutz et al. 1996); and secondly, the 3.3 μm PAH emission and 3.4 μm carbonaceous dust absorption features can be used to distinguish between the different energy sources of ULIRGs, without serious uncertainty in the continuum determination. Ground-based 3–4 μm spectroscopy is thus potentially a powerful tool to investigate the energetic importance of AGNs buried in the compact nuclei of ULIRGs, and to settle the controversy discussed above. Our particular interest is the search for buried AGNs in ULIRGs with cool far-infrared colors and/or non-Seyfert optical spectra, which are typically taken to be star-formation-dominated, based on *ISO* studies (Genzel & Cesarsky 2000; Taniguchi et al. 1999; Lutz, Veilleux, & Genzel 1999). UGC 5101 is such a ULIRG (Table 1 in this paper; Veilleux et al. 1999a), and has actually been diagnosed as being star-formation powered (Rigopoulou et al. 1999). In this Letter, we report on the results of 3–4 μm spectroscopic observations of UGC 5101, and the discovery of evidence for an energetically important buried AGN in this source. Throughout this paper, $H_0 = 75 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$ are adopted.

2. Observations and Data Analysis

We used the NSFCAM grism mode (Shure et al. 1994) to obtain a 3–4 μm spectrum of UGC 5101 at IRTF on Mauna Kea, Hawaii on the night of 2001 April 9 (UT). Sky conditions were photometric throughout the observations, and the seeing was measured to be $0''.7$ – $0''.8$ in full-width at half maximum. The detector was a 256×256 InSb array. The HKL grism and L blocker were used with the 4-pixel slit ($= 1''.2$). The resulting spectral resolution was ~ 150 at 3.5 μm .

The spectrum of UGC 5101 was obtained toward the flux peak at 3–4 μm . The position angle of the slit was 0 degree east of north. A standard telescope nodding technique with a throw of $12''$ was employed along the slit to subtract background emission. HR 4112 (F8V, $V=4.8$) was observed with an airmass difference of < 0.1 to correct for the transmission of the Earth’s atmosphere, and provide flux calibration.

Standard data analysis procedures were employed, using IRAF³. Initially, bad pixels were replaced with the interpolated signals of the surrounding pixels. Bias was subsequently subtracted from the obtained frames and the frames were divided by a spectroscopic flat frame. Finally the

³IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc. (AURA), under cooperative agreement with the National Science Foundation.

spectra of UGC 5101 and HR 4112 were extracted. Wavelength calibration was performed taking into account the wavelength-dependent transmission of the Earth’s atmosphere. Since we set each exposure time to 1.2 sec for the UGC 5101 observation, to reduce observing overheads, data at $> 4 \mu\text{m}$ were affected by the non-linear response of the detector. We excluded these data from our analysis. The UGC 5101 spectrum was then divided by the observed spectrum of HR 4112 and multiplied by the spectrum of a blackbody with a temperature appropriate to F8V stars (6200K). By adopting $L = 3.4$ for HR 4112 based on $V - L = 1.4$ (Tokunaga 2000), a flux-calibrated spectrum was produced.

3. Results

A flux-calibrated 3–4 μm spectrum of UGC 5101 is shown in Figure 1. At 3–4 μm , UGC 5101 was spatially very compact, and its spatial extent along the slit direction was indistinguishable from a point source. The spatially-unresolved red near-infrared nucleus reported by Scoville et al. (2000) and interpreted as a possible AGN by these authors is the probable origin of the bulk of the detected 3–4 μm continuum in our spectrum. Our spectrum gives a value of L (3.55 μm) of 10.1 mag, which is similar to the value of L' (3.7 μm) of 9.8 mag measured with a $5''$ aperture (Sanders et al. 1988a). Given that UGC 5101 has a red near-infrared color ($K - L' = 1.3$; Sanders et al. 1988a), so that $L - L'$ should be > 0 mag, our slit loss for the continuum emission is less than 0.3 mag. The dust continuum emission at 10–20 μm (Soifer et al. 2000), Pa α emission (Genzel et al. 1998, the top-middle of Fig 8), and 6–18 cm radio continuum emission (Sopp & Alexander 1991; Crawford et al. 1996) all show a centrally peaked morphology, with spatially-extended emission presumably originating in weakly obscured star-formation activity. Our $1''.2$ slit covers the bulk of this emission, so that most of the PAH emission should also be covered with our slit. Any missing PAH flux is very likely to be much smaller than the detected PAH flux, and will not seriously affect our quantitative discussion of the energy source of UGC 5101 (§ 4.1).

The spectrum in Fig. 1 shows a remarkable deviation at 3.3–3.7 μm from the smooth continuum emission. (A linear continuum level is shown as the solid line in Fig. 1.) We attribute the emission feature to 3.3 μm PAH emission, since the flux peak at 3.42 μm is consistent with the expected peak wavelength of the redshifted PAH emission ($3.29 \mu\text{m} \times 1.040$). Based on our adopted linear continuum level, the observed flux and rest-frame equivalent width of the 3.3 μm PAH emission are estimated to be $f_{\text{PAH}} = 1.3 \times 10^{-13} \text{ ergs s}^{-1} \text{ cm}^{-2}$ and $\text{EW}_{\text{PAH}} = 0.025 \mu\text{m}$, respectively.

At wavelengths longward of the PAH emission, signals at 3.5–3.7 μm are suppressed compared to the continuum level. The flux level suddenly and steeply decreases at $\sim 3.7 \mu\text{m}$, compared to the extrapolation from longer wavelengths, so that a strong absorption feature is undoubtedly present. We attribute this flux suppression to the 3.4 μm carbonaceous dust absorption (Pendleton et al. 1994), because the observed maximum of absorption at $\sim 3.53 \mu\text{m}$ is consistent with the redshifted absorption peak wavelength of the carbonaceous dust absorption feature ($3.4 \mu\text{m} \times 1.040$). Its observed optical depth is $\tau_{3.4}(\text{observed}) \sim 0.7$.

The 3.3 μm PAH emission profile almost fades out at 3.33–3.34 μm (Tokunaga et al. 1991; Imanishi & Dudley 2000), while the 3.4 μm carbonaceous dust absorption profile is just beginning at this wavelength range (Pendleton et al. 1994). At $z = 0.040$, this wavelength range, where neither emission nor absorption is particularly important, is redshifted to 3.46–3.47 μm . It will be seen that in Fig. 1 the adopted linear continuum intersects the observed data points at this wavelength, providing further evidence that the continuum determination is reasonable. An alternative hypothesis is represented by the curved continuum shown as the dashed line in Fig. 1. If this model were adopted, the observed flux and rest frame equivalent width of the 3.3 μm PAH emission would increase to $f_{\text{PAH}} = 2.0 \times 10^{-13} \text{ ergs s}^{-1} \text{ cm}^{-2}$ and $\text{EW}_{\text{PAH}} = 0.045 \mu\text{m}$, respectively, and the observed optical depth of the 3.4 μm dust absorption would decrease to $\tau_{3.4}(\text{observed}) \sim 0.6$. However, the level of this curved continuum seems to be too low; it requires that the absorption optical depth be close to zero at 3.48 μm or 3.35 μm in the rest frame, where the optical depth should in fact be significant (Pendleton et al. 1994). We thus believe that the actual PAH flux and equivalent width are lower than the values inferred using the curved continuum level; we can use these latter values as stringent upper limits. Considering the uncertainty of the continuum determination, we combine the values measured based on the two continuum levels and adopt $f_{\text{PAH}} = 1.6 \pm 0.3 \times 10^{-13} \text{ ergs s}^{-1} \text{ cm}^{-2}$, $\text{EW}_{\text{PAH}} = 0.035 \pm 0.010 \mu\text{m}$, and $\tau_{3.4}(\text{observed}) = 0.65 \pm 0.05$.

4. Discussion

4.1. A buried AGN in UGC 5101

The detection of 3.3 μm PAH emission indicates the presence of star-formation activity. However, its observed luminosity is $5.2 \pm 1.0 \times 10^{41} \text{ ergs s}^{-1}$, which yields an observed 3.3 μm PAH to far-infrared luminosity ratio (Table 1) of $\sim 1 \times 10^{-4}$, an order of magnitude smaller than is found in star-formation-dominated galaxies ($\sim 1 \times 10^{-3}$; Mouri et al. 1990). The detected, weakly-obscured star-formation activity thus contributes little to the huge far-infrared luminosity of UGC 5101 and a dominant energy source must be located behind the dust.

The presence of such a dust-obscured energy source is supported by the detection of strong 3.4 μm carbonaceous dust absorption. If the 3–4 μm continuum emission source behind the dust originated in star-formation activity, the rest-frame equivalent width of the 3.3 μm PAH emission should be similar to those of less obscured star-formation-dominated galaxies ($\sim 0.12 \mu\text{m}$; Imanishi & Dudley 2000), because both 3–4 μm continuum and 3.3 μm PAH emission fluxes would be attenuated similarly by dust extinction. The observed rest-frame equivalent width is, however, smaller by more than a factor of three. Consequently, the emission from behind the dust shows virtually no PAH emission, suggesting that the source is an AGN. This AGN activity dominates the observed 3–4 μm continuum flux, contributing $\sim 70\%$ of it, and is plausibly the energy source of UGC 5101’s far-infrared luminosity. If the dust obscuring such an energetically important AGN had a torus-like geometry, we would expect UGC 5101 to be optically classified as a Seyfert 2, but

in fact the non-Seyfert optical classification of UGC 5101 (Veilleux et al. 1999a) implies that the AGN is obscured by dust along all lines of sight (that is, a buried AGN).

After subtracting the contribution from star-formation to the observed 3–4 μm fluxes by using the spectral shape of the starburst galaxy NGC 253 in Imanishi & Dudley (2000), we estimate the intrinsic optical depth of the 3.4 μm carbonaceous dust absorption toward the buried AGN to be $\tau_{3.4}(\text{intrinsic}) \sim 0.8$, which yields dust extinction of $A_V > 100$ mag if a Galactic extinction curve is assumed (Pendleton et al. 1994). Assuming $A_{3-4\mu\text{m}} \sim 0.05 \times A_V$ (Lutz et al. 1996), we estimate the dereddened 3–4 μm dust emission luminosity powered by the buried AGN to be $\nu L_\nu \gtrsim 10^{45}$ ergs s $^{-1}$, which is comparable to the *observed* peak dust emission luminosity at 60 μm and 100 μm ($\nu L_\nu \sim 2 \times 10^{45}$ ergs s $^{-1}$, as calculated based on the *IRAS* fluxes shown in Table 1). In the case of a buried AGN, dust radiative transfer controls the temperature of the dust shell; the temperature of the dust decreases with increasing distance from the central AGN. The entire luminosity is transferred at each temperature. The present data thus provide evidence for an inner ~ 900 K, 3–4 μm continuum emitting dust shell of a luminosity similar to that of the observed outer ~ 40 K, 60–100 μm continuum emitting dust shell, as is expected from the buried AGN scenario.

For obscured AGNs, Alonso-Herrero, Ward, & Kotilainen (1997) found that the column density of X-ray absorbing gas, parameterized by N_{H} , relative to dust extinction towards the 3–4 μm continuum emitting region (A_V) is often higher by a large factor than the Galactic N_{H}/A_V ratio (1.8×10^{21} cm $^{-2}$ mag $^{-1}$; Predehl & Schmitt 1995). Dust obscuration toward the 3–4 μm continuum emitting region around the buried AGN in UGC 5101 is so high ($A_V > 100$ mag: this work) that N_{H} could easily exceed 10^{24} cm $^{-2}$, in which case direct 2–10 keV X-ray emission from the buried AGNs would be very strongly attenuated. The non-detection of 2–10 keV X-rays from the putative buried AGN with ASCA (Nakagawa et al. 1999) can thus be explained without difficulty.

The radio to far-infrared flux ratio of UGC 5101 is a factor of two larger than the typical values for star-formation-dominated galaxies at 1.5 GHz (Crawford et al. 1996), but inside their scatter at 151 MHz and 5 GHz (Cox et al. 1988; Sopp & Alexander 1991). Though the bulk of the radio emission is extended (Crawford et al. 1996; Sopp & Alexander 1991), the spatially-unresolved VLBI radio core, interpreted as an AGN by Smith et al. (1998), is luminous enough to account for the bolometric luminosity of UGC 5101 with AGN activity (Lonsdale et al. 1995).

4.2. Implications for optically LINER-type ULIRGs

Besides UGC 5101, studies in the thermal infrared have provided evidence for the presence of energetically dominant buried AGNs in two other optically non-Seyfert ULIRGs, IRAS 08572+3915 (Dudley & Wynn-Williams 1997; Imanishi & Dudley 2000) and IRAS F00183–7111 (Tran et al. 2001). The far-infrared emission properties of these three ULIRGs are summarized in Table 1. In the optical, all are classified as LINERs (Veilleux et al. 1999a; Lutz et al. 1999). For ULIRGs classified optically as LINERs, which constitute $\sim 40\%$ of ULIRGs (Veilleux et al. 1999a), two arguments

based on *ISO* studies have been made: (1) star-formation activity is energetically dominant, and (2) the LINER-type optical line emission is due to superwind-driven shocks caused by star-formation activity (Taniguchi et al. 1999; Lutz et al. 1999). The first argument, however, is not applicable to the three ULIRGs in Table 1. Then what is the origin of the LINER-type optical emission in these objects?

In buried AGNs, UV emission is blocked at the inner surface of the surrounding dust shell, but X-rays can penetrate deeply into the dust to produce X-ray dissociation regions (XDRs; Maloney, Hollenbach, & Tielens 1996). In XDRs, locally-generated UV photons, such as Lyman-Werner band H_2 emission and $\text{Ly}\alpha$ emission, can ionize some species with ionization potential less than 10–11 eV (notably carbon and iron), but the emitting volume is dominated by largely neutral gas, which causes XDRs essentially always to have LINER-type spectra (Maloney, in preparation). A buried AGN can thus explain the observed LINER-type optical emission.

However, in practical cases, some star-formation activity may very well be present at the surface of galaxies even when buried AGNs are energetically dominant, as we infer for UGC 5101. In such a case, due to the susceptibility of optical emission to dust extinction, the optical spectral diagnostics are likely to be dominated by this surface star-formation emission. The star-formation activity could be responsible for LINER-type optical emission via shocks, or might even alter the optical emission to a HII region type. It is important to note that narrow-line emission from high-excitation clouds characteristic of Seyfert galaxy spectra which are observed to be produced along the axes of nuclear dust tori would not contribute substantially at any wavelength in the case of buried AGN.

5. Summary

We have found evidence for the presence of an energetically dominant, buried AGN in the ULIRG UGC 5101, which is characterized by a cool far-infrared color and a LINER optical spectrum. This finding seems contrary to the currently established view, based on *ISO* data, that such objects are powered by star formation. Two other LINER ULIRGs also show evidence for energetically dominant buried AGNs, and we therefore conclude that at least some fraction of LINER ULIRGs are powered by buried AGNs.

We are grateful to J. Rayner and B. Golisch for their support before and during the observing run. Research in infrared astronomy at NRL is supported by the Office of Naval Research. PRM is supported by the NSF under grant AST-9900871. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

REFERENCES

- Alonso-Herrero, A., Ward, M. J., & Kotilainen, J. K. 1997, MNRAS, 288, 977
- Barger, A. J., Cowie, L. L., & Richards, E. A. 2000, AJ, 119, 2092
- Blain, A. W., Kneib, J. -P., Ivison, R. J., & Smail, I. 1999a, ApJ, 512, L87
- Blain, A. W., Smail, I., Ivison, R. J., & Kneib, J. -P. 1999b, MNRAS, 302, 632
- Cox, M. J., Eales, S. A. E., Alexander, P., & Fitt, A. J. 1988, MNRAS, 235, 1227
- Crawford, T., Marr, J., Partridge, B., & Strauss, M. A. 1996, ApJ, 460, 225
- Dudley, C. C. 1999, MNRAS, 307, 553
- Dudley, C. C., & Wynn-Williams, C. G. 1997, ApJ, 488, 720
- Fischer, J. 2000, astro-ph/0009395
- Genzel, R., & Cesarsky, C. J. 2000, ARA&A, 38, 761
- Genzel, R. et al. 1998, ApJ, 498, 579
- Imanishi, M., & Dudley, C. C. 2000, ApJ, 545, 701
- Lonsdale, C. J., Smith, H. E., & Lonsdale, C. J. 1995, ApJ, 438, 632
- Lutz, D. et al. 1996, A&A, 315, L269
- Lutz, D., Veilleux, S., & Genzel, R. 1999, ApJ, 517, L13
- Maloney, P., Hollenbach, D., & Tielens, A. G. G. M. 1996, ApJ, 466, 561
- Mouri, H., Kawara, K., Taniguchi, Y., & Nishida, M. 1990, ApJ, 356, L39
- Nakagawa, T., Kii, T., Fujimoto, R., Miyazaki, T., Inoue, H., Ogasaka, Y., Arnaud, K., & Kawabe, R. 1999, Galaxy Interactions at Low and High Redshift, (eds.) Barnes, J. E., and Sanders, D. B., IAU, 186, 341
- Pendleton, Y. J., Sandford, S. A., Allamandola, L. J., Tielens, A. G. G. M., & Sellgren, K. 1994, ApJ, 437, 683
- Predehl, P., & Schmitt, J. H. M. M. 1995, A&A, 293, 889
- Rigopoulou, D., Spoon, H. W. W., Genzel, R., Lutz, D., Moorwood, A. F. M., & Tran, Q. D. 1999, AJ, 118, 2625
- Roche, P. F., Aitken, D. K., Smith, C. H., & Ward, M. J. 1991, MNRAS, 248, 606

- Sanders, D. B., & Mirabel, I. F. 1996, *ARA&A*, 34, 749
- Sanders, D. B., Soifer, B. T., Elias, J. H., Madore, B. F., Matthews, K., Neugebauer, G., & Scoville, N. Z. 1988a, *ApJ*, 325, 74
- Sanders, D. B., Soifer, B. T., Elias, J. H., Neugebauer, G., & Matthews, K. 1988b, *ApJ*, 328, L35
- Scoville, N. Z. et al. 2000, *AJ*, 119, 991
- Shure, M. A., Toomey, D. W., Rayner, J. T., Onaka, P., & Denault, A. J. 1994, *Proc. SPIE*, 2198, 614
- Smith, H. E., Lonsdale, C. J., & Lonsdale, C. J. 1998, *ApJ*, 492, 137
- Soifer, B. T., Sanders, D. B., Madore, B. F., Neugebauer, G., Danielson, G. E., Elias, J. H., Lonsdale, C. J., & Rice, W. L. 1987, *ApJ*, 320, 238
- Soifer et al. 2000, *ApJ*, 119, 509
- Sopp, H. M., & Alexander, P. 1991, *MNRAS*, 251, 112
- Spoon, H. W. W., Keane, J. V., Tielens, A. G. G. M., Lutz, D., & Moorwood, A. F. M. 2001, *A&A*, 365, L353
- Taniguchi, Y., Yoshino, A., Ohyama, Y., & Nishiura, S. 1999, *ApJ*, 514, 660
- Tokunaga, A. T. 2000, in *Allen’s Astrophysical Quantities*, ed. A. N. Cox (4th ed: AIP Press: Springer), Chapter 7, p.143
- Tokunaga A. T., Sellgren K., Smith R. G., Nagata T., Sakata A., Nakada Y., 1991, *ApJ*, 380, 452
- Tran, Q. D. et al. 2001, *ApJ*, 552, 527
- Veilleux, S., Kim, D. -C., & Sanders, D. B. 1999a, *ApJ*, 522, 113
- Veilleux, S., Sanders, D. B., & Kim, D. -C. 1999b, *ApJ*, 522, 139

Table 1. Summary of UGC 5101, IRAS 08572+3915, and IRAS F00183–7111.

Object	Redshift	f_{25}^a [Jy]	f_{60}^a [Jy]	f_{100}^a [Jy]	$\log L_{\text{FIR}}^b$ [ergs s $^{-1}$]	f_{25}/f_{60}^c
UGC 5101	0.040	1.03	11.54	20.23	45.61	0.09 (cool)
IRAS 08572+3915	0.058	1.70	7.43	4.59	45.62	0.23 (warm)
IRAS F00183–7111	0.327	0.13	1.20	1.19	46.51	0.11 (cool)

^a f_{25} , f_{60} , and f_{100} are *IRAS FSC* fluxes at $25\mu\text{m}$, $60\mu\text{m}$, and $100\mu\text{m}$, respectively.

^bLogarithm of far-infrared ($40\text{--}500\ \mu\text{m}$) luminosity in ergs s $^{-1}$ calculated with $L_{\text{FIR}} = 1.4 \times 2.1 \times 10^{39} \times D(\text{Mpc})^2 \times (2.58 \times f_{60} + f_{100})$ ergs s $^{-1}$ (Sanders & Mirabel 1996).

^c $f_{25}/f_{60} < (>) 0.2$ are called cool (warm) (Sanders et al. 1988b).

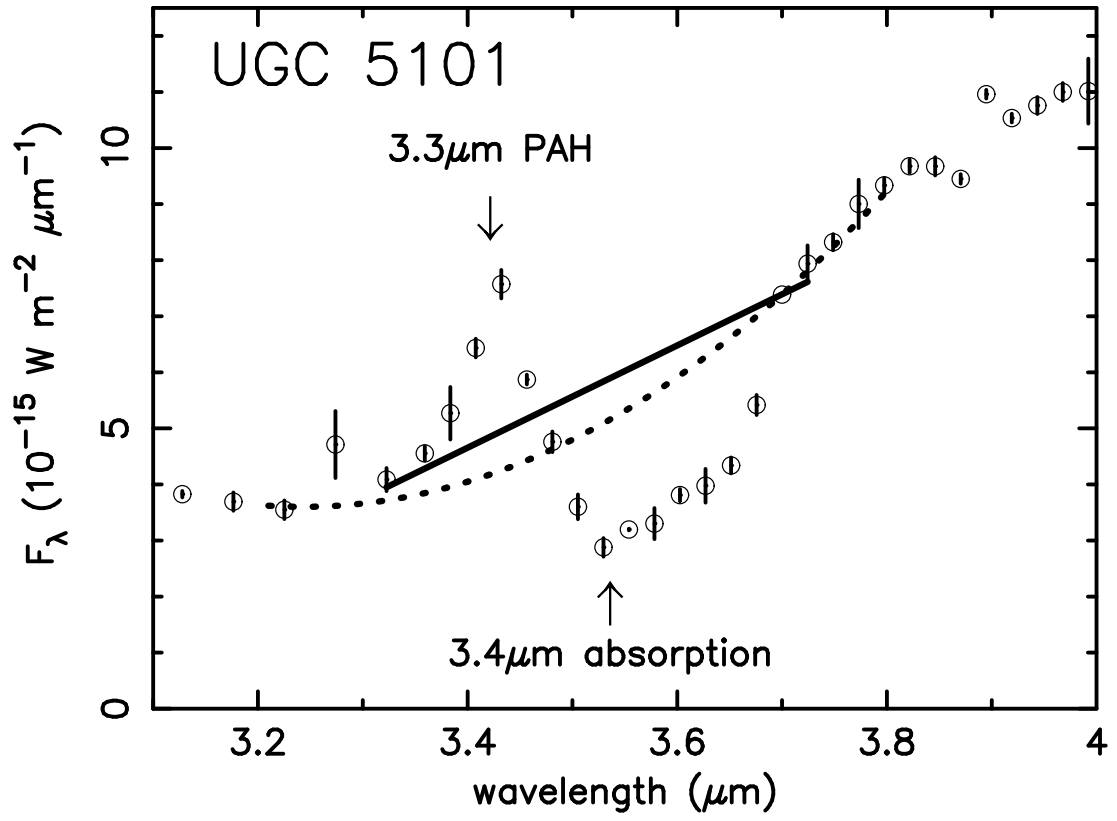


Fig. 1.— 3–4 μm spectrum of UGC 5101. The abscissa and ordinate are the observed wavelength and the flux in F_λ , respectively. The solid and dashed lines are possible continuum levels (see § 3).